

## O(He) STARS

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95% of all stars end their lives as white dwarfs. About 20% of the hot post-AGB stars are hydrogen deficient. Most of these are the result of a (very) late helium-shell flash, but the evolutionary status of a fraction of about 10 – 20% of the hottest hydrogen-deficient stars, namely the O(He) stars, is as yet unexplained.

For the evolution of O(He) stars there are two scenarios: They could be the long-sought hot successors of RCrB stars, which have not been identified up to now. If this turns out to be true, then a third post-AGB evolutionary sequence is revealed, which is probably the result of a double-degenerate merging process. An alternative explanation might be that O(He) stars are post early-AGB stars. These depart from the AGB just before they experience their first thermal pulse which will then occur as a late thermal pulse (LTP). This would be a link to the low-mass He-enriched PG 1159 stars.